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Radio and Gamma-ray Properties of Extragalactic Jets from the TANAMI Sample

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ABSTRACT

Using high-resolution radio imaging with VLBI techniques, the TANAMI program has been observing the parsec-scale radio jets of southern (declination south of -30°) γ -ray bright AGN simultaneously with *Fermi*/LAT monitoring of their γ -ray emission. We present the radio and γ -ray properties of the TANAMI sources based on one year of contemporaneous TANAMI and *Fermi*/LAT data. A large fraction (72%) of the TANAMI sample can be associated with bright γ -ray sources for this time range. Association rates differ for different optical classes with all BL Lacs, 76% of quasars and just 17% of galaxies detected by the LAT. Upper limits were established on the γ -ray flux from TANAMI sources not detected by LAT. This analysis led to the identification of three new *Fermi* sources whose detection was later confirmed. The γ -ray and radio luminosities are related by $L_\gamma \propto L_r^{0.89 \pm 0.04}$. The brightness temperatures of the radio cores increase with the average γ -ray luminosity, and the presence of brightness temperatures above the inverse Compton limit implies strong Doppler boosting in those sources. The undetected sources have lower γ /radio luminosity ratios and lower contemporaneous brightness temperatures. Unless the *Fermi*/LAT-undetected blazars are strongly γ -ray-fainter than the *Fermi*/LAT-detected ones, their γ -ray luminosity should not be significantly lower than the upper limits calculated here.

Key words. Galaxies: jets – Galaxies: nuclei – Galaxies:quasars: individual – Gamma rays: galaxies – Radio continuum: galaxies –

1. Introduction

Blazars are a subset of active galactic nuclei (AGN). They are very luminous, strongly variable, and show strong polarized emission (Urry & Padovani, 1995). These properties can be explained by emission from collimated jets consisting of charged particles moving at relativistic velocities that are oriented at a small angle to the line of sight and thus Doppler boosted (Blandford & Rees, 1978; Maraschi et al., 1992). This explanation is confirmed by the fact that blazars typically exhibit apparent superluminal motion in the inner radio jet (see, e.g., Lister et al., 2013, for an extensive study). While detailed understanding of their emission processes is still a work in progress (e.g., Ghisellini et al., 2009; Ghisellini et al., 2010; Tavecchio et al., 2010, and references therein), a close link between radio and high-energy emission from blazars is clear (Fossati et al., 1998; Kovalev et al., 2009; Ackermann et al., 2011a).

Detailed studies of AGN in the MeV-to-GeV energy regime started with the EGRET detector aboard the *Compton Gamma Ray Observatory* (CGRO; Thompson et al., 1993), which found that many blazars are strong γ -ray emitters (Hartman et al., 1992; Mattox et al., 1996; Bloom, 2008, and references therein). Radio observations, particularly using Very Long Baseline Interferometry (VLBI), were immediately recognized as a particularly useful tool to understand high-energy emission from blazars. VLBI provides information at the highest possible resolution and is the only way to measure kinematics of a blazar jet, including many of the parameters that are essential in-

puts to models that seek to explain blazar emission processes (e.g., Cohen et al., 2007). Early studies using EGRET and VLBI gave astronomers the first good glimpse of the high-energy blazar emission and its connection to radio properties, e.g., compared to non-detections, EGRET detections had a higher radio flux density and variability (Impey, 1996; Tingay et al., 2003), higher brightness temperatures (Moellenbrock et al., 1996), more strongly polarized jets (Lister & Homan, 2005) and larger than average opening angles (Taylor et al., 2007).

Despite the insights provided into high-energy blazar emission, EGRET data were limited in many ways, such as by the difficulty of precise determination of source positions, the very non-uniform sky coverage, poor temporal sampling, and the limited sensitivity. For that reason many results were tentative, incomplete, or even inconsistent. Most obviously, many of the most radio luminous and compact blazars were not detected. A greatly enhanced successor to EGRET, the Large Area Telescope (LAT) on board the *Fermi Gamma-ray Space Telescope* (Atwood et al., 2009), was launched on 2008 June 11 and commenced regular observations two months later. *Fermi*/LAT is a pair conversion detector of γ -rays with energies in the range ~ 20 MeV to >300 GeV. The sensitivity of LAT is more than an order of magnitude higher than that of EGRET. In LAT's sky-survey mode the entire sky is scanned every 3 hours, and fairly uniform exposure is obtained within two months. One of LAT's major scientific goals is to observe the γ -ray activity of AGN: to detect, monitor, and characterize rapidly variable flaring sources. Most of the EGRET detections have been

confirmed by *Fermi*. Detailed discussions of AGN detected with *Fermi*/LAT are given in “The First Catalog of Active Galactic Nuclei Detected by the *Fermi* Large Area Telescope” (1LAC; Abdo et al., 2010d), and its second (2LAC; Ackermann et al., 2011b) and third revisions (3LAC; Ackermann et al., 2015).

The TANAMI (Tracking Active Galactic Nuclei with Austral Milliarsecond Interferometry; Ojha et al., 2010) program is a VLBI monitoring program targeting AGN jets south of -30° declination. Observations are made at two radio frequencies (8 and 22 GHz) approximately every two months using the telescopes of the Australian Long Baseline Array (LBA; e.g. Ojha et al., 2004) in combination with telescopes in Australia (NASA’s Tidbinbilla facility), South Africa, Antarctica, Chile, and New Zealand. The array has been further expanded by the inclusion of one of the antennas of the ASKAP (Australian Square Kilometre Array Pathfinder) array and the Warkworth antenna in New Zealand (see, e.g., Tzioumis et al., 2010) and the new AuScope antennas at Yarragadee (Western Australia) and Katherine (Northern Territory) (see, e.g., Lovell et al., 2013). The dual frequency nature of the VLBI observations yields spectral index maps of parsec-scale jet features. The multi-epoch monitoring allows determination of jet parameters such as jet speed and collimation angles via tracking of individual jet components. As these critical parameters cannot be determined by any other observational technique, the highest possible spatial resolution provided by VLBI is uniquely important to understanding high-energy emission processes in jets.

Complementary to correlations between the total unresolved radio and the gamma-ray flux (e.g., Ackermann et al., 2011a), observations by TANAMI and other groups have shown that the γ -ray brightness of AGN as seen by LAT is correlated with VLBI jet properties, such as the opening angle or jet speed, and found other connections, e.g. between γ -ray loudness and synchrotron peak frequency (Lister et al., 2009; Kovalev et al., 2009; Ojha et al., 2010; Lister et al., 2011; Linford et al., 2012). We present the 8.4 GHz radio and the γ -ray properties of the TANAMI sample, as obtained from an analysis of the first 11 months of *Fermi*/LAT observations, which is the period of time used for the First LAT Catalog (1FGL; Abdo et al., 2010c). The 1FGL data set covers the period from 2008 August 4 to 2009 July 4. Our analysis is based on the 1FGL period, because the TANAMI sample has been selected to include southern AGN detected with *Fermi*/LAT in this period of time (details of the sample selection are provided in Sect. 2.1). We use only γ -ray data covering the same period of time as the radio data. Due to source variability, the usage of non-simultaneous data would wash out the signal from these correlation studies between the two energy regimes; Pushkarev et al. (2010) find a radio/ γ -ray delay on the order of months for the best correlation. At higher radio frequencies shorter time delays are observed. By studying variability of individual sources León-Tavares et al. (2011) found that the brightest γ -ray emission (seen with *Fermi*/LAT) occurs in the rising phase of millimeter flares.

We present our analysis approach in Sect. 2. Results of the analysis are presented in Sect. 3, including the detection statistics and possible non-1FGL γ -ray counterparts of TANAMI sources. Section 4 is devoted to a discussion of some individual sources. We discuss our results in Sect. 5 and end with our conclusions in Sect. 6.

2. Analysis

2.1. The TANAMI sample

The TANAMI sample, which is analyzed here, contains 75 AGN. It is defined as a combined radio and γ -ray selected sample that includes most radio-loud extragalactic jets south of $\delta = -30^\circ$ that have either been detected at γ -ray energies or are considered candidate γ -ray sources.

The sample of 75 AGN includes the initial TANAMI sample of 43 southern sources. The latter consists of a radio-selected flux-density-limited sub-sample (with $S_{5\text{GHz}} > 2\text{ Jy}$), a γ -ray selected sub-sample of known and candidate γ -ray loud jets based on results of *CGRO*/EGRET, and sources from special classes, such as intra-day variable (IDV) and GHz peaked spectrum (GPS) sources. The sample selection and the radio properties of the first 43 sources are discussed by Ojha et al. (2010). During the first months of *Fermi* operation, several southern AGN that are candidate counterparts for LAT sources were added to the initial sample of 43 AGN, resulting in a total of 75 objects in the TANAMI sample. This sample does not include all 101 AGN in this declination range that have been detected with *Fermi*/LAT in the 1FGL time range.

Based on their optical properties, the TANAMI sample is classified into quasars (Q), BL Lac objects (B), and galaxies (G). These classifications are based on Véron-Cetty & Véron (2006) and Shaw et al. (2012, 2013), with the exception of a few cases where we have updated the source classification based on newer references (see Sect. 4). In total the sample contains 38 quasars, 16 BL Lac objects, and 11 radio galaxies. For the remaining 10 sources no classifications were available, because most of them are faint in the optical.

2.2. γ -ray counterparts and upper limits

In order to study the radio/ γ -ray connection, we searched for sources in the TANAMI sample with γ -ray counterparts in the 1FGL catalog. The associations are based on positional coincidence of the radio and γ -ray source positions. The associations used are consistent with those in the 1FGL catalog and its corresponding AGN catalog, 1LAC (Abdo et al., 2010d).

For TANAMI sources without a counterpart in the 1FGL catalog we calculated upper limits on the γ -ray flux at the corresponding radio positions. The upper limits and the test statistic (TS) of these sources were obtained by a maximum likelihood analysis (Cash, 1979; Mattox et al., 1996), where $\sqrt{\text{TS}}$ is comparable to the significance in σ . We used only photons in the “Source” class of P7_V6 events with energies in the range 100 MeV–100 GeV for the calculation of upper limits. To minimize contamination from Earth’s limb γ -rays, photons with zenith angles greater than 100° were removed. The standard *Fermi*/LAT *ScienceTools* software package¹ (version v9r23p1) was used with the “P7SOURCE_V6” set of instrument response functions². The flux, photon index, and TS of each source were determined by analyzing a Region of Interest (RoI) of 10° in radius centered at the radio position. We modeled the LAT point sources with individual power-law spectra (photon flux $dN/dE = K(E/E_0)^{-\Gamma}$). The Galactic diffuse background (gal_2yearp7v6_v0) and the isotropic background

¹ <http://fermi.gsfc.nasa.gov/ssc/data/analysis/documentation/Cicerone/>

² The 1FGL analysis was done using the P6_V3 events and response functions. We used the 1FGL results because the fluxes and spectral indices did not change significantly with the P7 analysis, as shown by the 2FGL catalog (Nolan et al., 2012).

(iso_p7v6source) used in the RoI model, including the γ -ray diffuse and residual background of misclassified cosmic rays, are the recommended versions released and described in more detail in the documentation available at the *Fermi* Science Support Center³. The xml source model for the analyzed region has been created using the 1FGL catalog and a modified version of the tool `make2FGLxml.py` contributed by T. Johnson, where a source has been added at the radio position for the upper limit calculation. The tool creates a source model including all sources in the RoI as well as sources that are close enough to contribute photons in the RoI due to the instrument response function of *Fermi*/LAT. Model parameters from the 1FGL catalog are used as default values for every source in the model. We fixed the model parameters of sources outside the RoI to their catalog values, while parameters of sources inside the RoI were variable in the modelling process.

Uncertainties in the LAT effective area represent the major source of systematic error in these results. These uncertainties in the effective area for the IRFs were evaluated by Ackermann et al. (2012) as 10% at 100 MeV, 5% at 560 MeV, and 10% above 10 GeV, linearly varying with the logarithm of energy between those values. The statistical uncertainties exceed these values in all cases. The reported errors on spectral parameters are 1σ uncertainties and statistical only. Following Ackermann et al. (2012), we estimate that the systematic uncertainties are comparable or smaller, $\sim 8\%$ for the fluxes and ~ 0.1 in photon indices.

In a few cases this LAT analysis of the unassociated AGN from the TANAMI sample using the enhanced sensitivity of the P7 data gave detections ($TS \geq 25$) instead of upper limits (see Sect. 3.2 for more information). For the non-detected sources upper limits on the γ -ray flux have been obtained by fixing the spectral index to $\Gamma = 2.4$ and increasing the flux until a ΔTS of 2.71 was reached, which yields an upper limit at the 90% confidence limit. As this method underestimates the upper limit for sources with $TS < 1$, the Bayesian method (Helene, 1983) has been applied in these cases (for more information see Sect. 4.4 of Abdo et al., 2010c).

Throughout this work, the γ -ray band 100 MeV–100 GeV is used. For comparisons of γ -ray properties obtained in analyses done here with corresponding quantities reported in the LAT catalogs (Nolan et al., 2012), values such as spectral index and energy flux could be taken directly from these catalogs because the energy ranges analyzed were the same. As the integrated flux in the 100 MeV–100 GeV band is not given directly in the published catalogs, we calculated this value using the flux density at the pivot energy and the spectral index (see, e.g., Abdo et al., 2010c, for more information). The flux uncertainty in this band has been obtained using the uncertainties of the spectral index and the flux density.

2.3. Radio Analysis

The VLBI radio analysis of the TANAMI sources follows Ojha et al. (2010), who published results for the initial 43 sources in the sample. For each source we used the first radio epoch with an observation date within the *Fermi* 1FGL period as being representative of the source's radio flux during the 1FGL period.

We determined brightness temperatures for the radio core at 8.4 GHz in the source frame as described by Ojha et al. (2010),

$$T_B = \frac{2 \ln 2}{\pi k_B} \frac{S_{\text{core}} \lambda^2 (1+z)}{\theta_{\text{maj}} \theta_{\text{min}}} \quad (1)$$

where S_{core} , θ_{maj} , and θ_{min} are the flux density (in Janskys), the semimajor, and semiminor axis of a two dimensional Gaussian model for the core in the radio image (in milliarcseconds), k_B is the Boltzmann constant, z the redshift of the source, and λ the observing frequency. If the size of the fitted model component for the core emission falls below the resolution limit, we calculated lower limits for the brightness temperature following Kovalev et al. (2005). The sources for which only a lower limit on T_B can be given are PKS 0717–432, PKS 0812–736, and PKS 1606–667 (see Sect. 3.2).

For a few sources in the sample, no radio properties are presented here, because either they were not observed in the 1FGL period or no radio core could be identified in their VLBI image due to an irregular morphology. Although these sources do not contribute to the study of radio and γ -ray emission, we show their γ -ray properties for completeness.

2.4. Gamma-ray Luminosities

Assuming the sources have a power-law photon flux spectrum of the form

$$N_{\text{ph}}(E) = S_{\text{ph}} \frac{1-\Gamma}{E_0} \left(\left(\frac{E_{\text{max}}}{E_0} \right)^{1-\Gamma} - \left(\frac{E_{\text{min}}}{E_0} \right)^{1-\Gamma} \right)^{-1} \left(\frac{E}{E_0} \right)^{-\Gamma} \quad (2)$$

where S_{ph} is the measured photon flux in the energy band from E_{min} to E_{max} (E_0 is only a reference energy providing a dimensionless base for the non-integer exponent), and where the photon index $\Gamma \neq 1$. The energy flux in that band is given by

$$S_E = S_{\text{ph}} E_0 \frac{1-\Gamma}{2-\Gamma} \frac{(E_{\text{max}}/E_0)^{2-\Gamma} - (E_{\text{min}}/E_0)^{2-\Gamma}}{(E_{\text{max}}/E_0)^{1-\Gamma} - (E_{\text{min}}/E_0)^{1-\Gamma}} \quad (3)$$

for $\Gamma \neq 2$.

As the sources in the sample are located at different distances, we corrected the measured luminosities using a K-correction following Ghisellini et al. (2009), i.e.,

$$L_E = 4\pi d_L^2 \frac{S_E}{(1+z)^{2-\Gamma}} \quad (4)$$

where z is the redshift of the source, S_E the energy flux, and Γ the photon spectral index in the γ -ray band. The luminosity distance, d_L , was calculated assuming a flat universe with $\Omega_M = 0.27$, $\Omega_\Lambda = 0.73$, and $H_0 = 71.0 \text{ km s}^{-1} \text{ Mpc}^{-1}$.

3. Results

3.1. Detection Statistics

The analysis of LAT data revealed different results for different optical classes of AGN. A total of 54 of the 75 AGN from the TANAMI sample can be associated with γ -ray sources from the 1FGL catalog. Table 1 lists the associated AGN. All of the BL Lac objects (16/16) in the sample were detected in the γ -ray regime, and 29 out of 38 quasars were detected, but only 2 out of 11 radio galaxies have strong enough γ -ray emission to be detected in the analysed 11 months of LAT data. The low detection fraction for the radio galaxies is consistent with jet inclination

³ <http://fermi.gsfc.nasa.gov/ssc/data/access/lat/BackgroundModels.html>

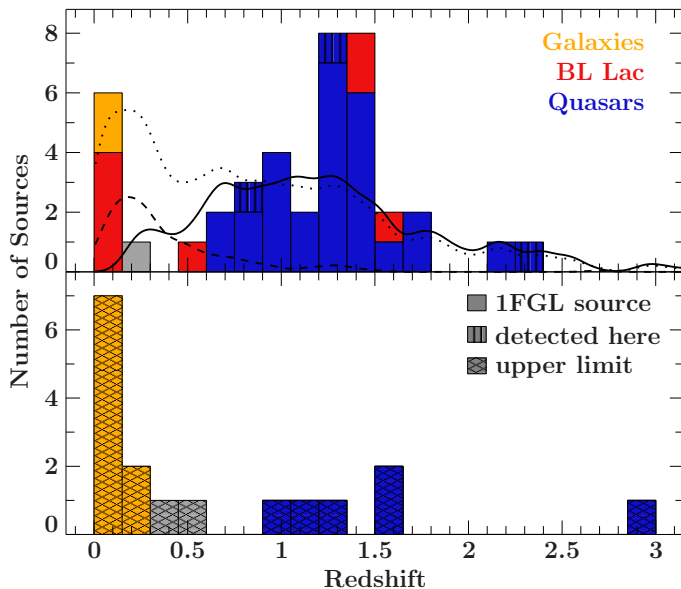


Fig. 1. Redshift distribution of LAT detected (top panel) and non-detected sources (bottom panel). The lines in the top panel represent the redshift distributions of all *Fermi*/LAT-detected AGN (dotted line), of BL Lac objects (dashed line), and quasars (solid line) in the 2LAC, which are shown by a kernel density estimation (KDE) scaled to the corresponding numbers of TANAMI sources.

effects and relativistic beaming as predicted by AGN unification (Urry & Padovani, 1995). Out of the 10 unclassified sources, 7 are detected with LAT. We note that the fraction of γ -ray detections of certain source classes is biased by our inhomogeneous sample selection (see Sect. 2.1), e.g., sources have been added due to a LAT detection. The added sources include many unclassified objects, whereas the radio galaxies were all included in the initial TANAMI sample.

The two *Fermi*/LAT-detected radio galaxies are PKS 1322–428 (Cen A) and PKS 0521–365 (ESO 362-G021). Cen A, which is the closest AGN, has the lowest γ -ray luminosity of all associated sources. Gamma-ray emission from its central region and from the giant radio lobes is observed (Abdo et al., 2010a,b). Due to Cen A’s proximity, properties of its jet can be studied with exceptionally high resolution at sub-parsec scales using radio VLBI (see, e.g., Tingay et al., 2001a; Müller et al., 2011, 2014). The other detected radio galaxy, PKS 0521–365, has been suggested to be a BL Lac object based on the properties of its nucleus (see, e.g., Danziger et al., 1979). It was considered as an example of a misaligned radio galaxy with an innermost jet on mas-scales oriented close to the line of sight. Interestingly, Tingay & Edwards (2002) find that this source is likely not “strongly affected by relativistic boosting”. Further TANAMI radio observations will help to clarify the nature of this source by obtaining detailed properties of its inner jet.

Figure 1 shows the distribution of redshifts of *Fermi*/LAT γ -ray detected and undetected sources (those detected in 1FGL and two detections reported in Sect. 3.2. The third detection does not have a measured redshift). No statistical differences between the redshift distributions of *Fermi*/LAT-detected and undetected sources of each class are found. A Kolmogorov-Smirnov (KS) two sample test does not indicate a significant difference between the redshift distribution of *Fermi*/LAT-detected quasars and that of non-detected quasars. The same applies to the radio

galaxies. However, it has to be noted that for the quasars, as well as for the radio galaxies, one of the compared distributions includes only a small number of elements. All radio galaxies in the sample are at low redshifts, and the majority of them remain undetected. The comparison of redshift distributions cannot be done with BL Lac objects, because all sources of this type in the TANAMI sample are detected with *Fermi*/LAT. Comparing the redshift distribution of the *Fermi*/LAT-detected AGN in the TANAMI sample with that of all *Fermi*/LAT-detected AGN given in 2LAC indicates slight differences. Relative to 2LAC, the TANAMI sample contains fewer sources in the moderate redshift regime of around 0.2–0.8. This difference is caused by the selection of the TANAMI sample. Contrary to the 2LAC “Clean Sample” (AGN with Galactic latitude $|b| > 10^\circ$), which includes 395 BL Lac objects and 310 quasars, the fraction of quasars is larger in the TANAMI sample with 38 quasars but only 16 BL Lac objects. While there is no obvious difference between the redshift distribution of quasars in TANAMI and 2LAC, the flux-limited sample selection of TANAMI seems to favor the nearby BL Lac objects. The distributions of galaxies in 2LAC are not shown separately, due to their low number and only small redshifts.

3.2. Possible Non-1FGL Gamma-ray Counterparts

For sample sources without 1FGL γ -ray counterparts, we calculated upper limits on the γ -ray flux as described in Sect. 2.2, using the first 11 months of *Fermi*/LAT data. For some radio sources, this analysis revealed significant γ -ray emission at the radio position. We modeled sources with $TS \geq 25$ (Table 2) and calculated upper limits for less significant sources (Table 3). The detection of additional γ -ray sources not included in the 1FGL catalog benefited from our usage of the Galactic and extragalactic background models obtained during the first two years of *Fermi* operations, as well as the improved instrument response functions.

Two quasars (PKS 2149–306 and PKS 2326–477) and one source with unknown optical counterpart (PKS 1505–495) were detected with $TS \geq 25$, i.e., they met the detection threshold for the LAT source catalogs. All three of these detections confirm analysis by the *Fermi*/LAT team in the second year *Fermi*/LAT-catalog (2FGL; Nolan et al., 2012) and its AGN counterpart, 2LAC (Ackermann et al., 2011b).

3.3. Gamma-ray properties of the TANAMI sample

Figure 2 shows the γ -ray flux distribution of the sources in the TANAMI sample. The large number of unclassified sources mostly results from the addition of new *Fermi* detections that are proportionally fainter and thus less well studied at other frequencies. The flux and the spectral index are averaged over the first 11 months of *Fermi* science operations. Given the different source distances it is in general difficult to discern any clear connection between source type and γ -ray flux. The flux distributions of the AGN in 2LAC and in the TANAMI sample are clearly different, as the sample selection led to a much higher fraction of bright γ -ray sources.

The distribution of spectral indices is shown in Fig. 3. Typical uncertainties are in the range of ± 0.1 (Table 1). There is an indication that BL Lacs tend to have, on average, harder spectra than quasars, which is consistent with earlier *Fermi* results (Abdo et al., 2010d, their Fig. 12). With a KS two sample test we obtain a probability of 0.1% that the γ -ray spectral indices

Table 1. Properties of TANAMI sources associated with IFGL sources. Luminosities are only available for sources with measured redshift. For a few AGN no radio core flux is given, because either their morphology did not allow a unique identification of a core or they have not been observed within the IFGL time range.

Name ^a	Source	Common name	IFGL Association	TS	Flux (0.1–100 GeV) [10 ⁻⁸ ph cm ⁻² s ⁻¹]	LAT Γ	L (0.1–100 GeV) [erg s ⁻¹]	S_{core} (8.4 GHz) [Jy]	T_B^b [K]	ID ^c	z
0047–579	J0049–5738	[HB89] 0047–579	1FGL J0049.8–5738	80	2.3 ± 0.7	2.42 ± 0.17	(3.8 ± 1.5) × 10 ⁴⁷	1.35 ± 0.21	4.8 × 10 ¹¹	Q	1.797
0055–328	J0058–3234	PKS 0055–328	1FGL J0058.4–3235	61	1.6 ± 0.5	2.31 ± 0.16	(1.1 ± 0.5) × 10 ⁴⁷	0.08 ± 0.02	1.4 × 10 ¹⁰	B	1.370
0208–512	J0210–5101	[HB89] 0208–512	1FGL J0210.6–5101	1175	14.6 ± 1.0	2.37 ± 0.05	(4.7 ± 0.4) × 10 ⁴⁷	1.39 ± 0.21	2.3 × 10 ¹¹	Q	0.999
0227–369	J0229–3643	PKS 0227–369	1FGL J0229.3–3644	376	8.0 ± 0.8	2.60 ± 0.08	(2.0 ± 0.3) × 10 ⁴⁸	0.57 ± 0.09	5.1 × 10 ¹¹	Q	2.115
0244–470	J0246–4651	PKS 0244–470	1FGL J0245.9–4652	571	9.4 ± 0.8	2.52 ± 0.07	(6.9 ± 0.7) × 10 ⁴⁷	0.26 ± 0.04	1.9 × 10 ¹¹	Q	1.385
0302–623	J0303–6211	PKS 0302–623	1FGL J0303.4–6209	117	5.1 ± 1.1	2.59 ± 0.13	(3.7 ± 0.9) × 10 ⁴⁷	1.04 ± 0.16	1.5 × 10 ¹¹	Q	1.351
0308–611	J0309–6058	PKS 0308–611	1FGL J0310.1–6058	104	4.2 ± 1.1	2.53 ± 0.15	(3.0 ± 1.1) × 10 ⁴⁷	0.81 ± 0.13	2.2 × 10 ¹¹	Q	1.480
0332–403	J0334–4008	[HB89] 0332–403	1FGL J0334.2–4010	675	8.2 ± 0.7	2.35 ± 0.06	(6.0 ± 0.7) × 10 ⁴⁷	0.95 ± 0.15	1.8 × 10 ¹¹	B	1.351
0332–376	J0334–3725	PMN J0334–3725	1FGL J0334.4–3727	226	2.8 ± 0.5	2.10 ± 0.09		0.32 ± 0.05	2.5 × 10 ¹¹ *	B	>0.39
0402–362	J0403–3605	PKS 0402–362	1FGL J0403.9–3603	685	10.9 ± 0.9	2.56 ± 0.06	(8.6 ± 0.8) × 10 ⁴⁷	0.54 ± 0.09	1.6 × 10 ¹¹	Q	1.417
0405–385	J0406–3826	[HB89] 0405–385	1FGL J0407.4–3827	339	6.1 ± 0.8	2.40 ± 0.08	(3.8 ± 0.6) × 10 ⁴⁷	1.38 ± 0.21	1.4 × 10 ¹¹	Q	1.285
0412–536	J0413–5332	PMN J0413–5332	1FGL J0413.4–5334	352	7.4 ± 0.9	2.55 ± 0.09	(2.6 ± 0.4) × 10 ⁴⁷			Q	1.024
0426–380	J0428–3756	PKS 0426–380	1FGL J0428.6–3756	7750	31.5 ± 0.9	2.13 ± 0.02	(1.7 ± 0.1) × 10 ⁴⁸	1.78 ± 0.27	1.6 × 10 ¹²	Q	1.110
0447–439	J0449–4350	PKS 0447–439	1FGL J0449.5–4350	2221	9.7 ± 0.7	1.95 ± 0.04		0.09 ± 0.02	4.7 × 10 ¹⁰ *	B	
0454–463	J0455–4615	[HB89] 0454–463	1FGL J0455.6–4618	214	6.2 ± 0.9	2.57 ± 0.10	(1.3 ± 0.2) × 10 ⁴⁷	3.24 ± 0.49	4.2 × 10 ¹¹	Q	0.853
0506–612	J0506–6109	[HB89] 0506–612	1FGL J0507.3–6103	162	7.6 ± 1.3	2.73 ± 0.12	(2.7 ± 0.6) × 10 ⁴⁷	1.12 ± 0.17	3.3 × 10 ¹¹	Q	1.093
0516–621	J0516–6207	PKS 0516–621	1FGL J0516.7–6207	328	5.8 ± 1.0	2.28 ± 0.10	(4.6 ± 0.8) × 10 ⁴⁷	0.94 ± 0.15	3.6 × 10 ¹¹	Q	1.300
0521–365	J0522–3627	ESO 362- G 021	1FGL J0522.8–3632	605	11.5 ± 0.9	2.60 ± 0.07	(3.2 ± 0.3) × 10 ⁴⁴	1.52 ± 0.23	2.9 × 10 ¹¹	G	0.055
0524–485	J0526–4830	PKS 0524–485	1FGL J0526.3–4829	137	3.6 ± 0.7	2.37 ± 0.12	(2.4 ± 0.6) × 10 ⁴⁷	0.57 ± 0.09	2.2 × 10 ¹¹	Q	1.300
0537–441	J0538–4405	[HB89] 0537–441	1FGL J0538.8–4404	6494	37.8 ± 1.1	2.27 ± 0.03	(9.8 ± 0.4) × 10 ⁴⁷	4.71 ± 0.71	5.6 × 10 ¹²	Q	0.894
0625–354	J0627–3529	PKS 0625–35	1FGL J0627.3–3530	50	0.7 ± 0.4	1.86 ± 0.18	(6.7 ± 3.0) × 10 ⁴³	0.35 ± 0.06	1.2 × 10 ¹¹	B	0.055
0637–752	J0635–7516	[HB89] 0637–752	1FGL J0636.1–7521	110	5.1 ± 1.1	2.49 ± 0.13	(4.2 ± 1.1) × 10 ⁴⁶	3.02 ± 0.46	2.0 × 10 ¹¹	Q	0.653
0700–661	J0700–6610	PKS 0700–661	1FGL J0700.4–6611	486	6.0 ± 0.8	2.15 ± 0.07		0.14 ± 0.03	2.9 × 10 ¹⁰ *	B	>0.39
0717–432	J0718–4319	PMN J0718–4319	1FGL J0718.7–4320	248	1.9 ± 0.4	1.83 ± 0.09		0.03 ± 0.01	>2.6 × 10 ¹¹ *		
0736–770	J0734–7711	PKS 0736–770	1FGL J0734.1–7715	88	7.2 ± 1.3	2.75 ± 0.13				Q	
0745–330	J0747–3310	PKS 0745–330	1FGL J0747.4–3303	130	8.9 ± 1.4	2.38 ± 0.08					
0812–736	J0811–7530	PMN J0811–7530	1FGL J0811.1–7527	151	1.4 ± 0.4	1.80 ± 0.12		0.04 ± 0.01	>5.1 × 10 ¹¹ *		
1057–797	J1058–8003	PKS 1057–79	1FGL J1058.1–8006	154	6.3 ± 1.1	2.45 ± 0.11	(4.5 ± 0.9) × 10 ⁴⁶	1.93 ± 0.29	4.7 × 10 ¹¹	B	0.581
1101–536	J1103–5357	PKS 1101–536	1FGL J1103.9–5355	396	6.7 ± 0.9	2.05 ± 0.06		0.39 ± 0.06	6.2 × 10 ¹⁰ *	B	>0.15
1144–379	J1147–3812	[HB89] 1144–379	1FGL J1146.9–3812	221	4.9 ± 0.8	2.31 ± 0.09	(1.8 ± 0.3) × 10 ⁴⁷	1.91 ± 0.29	5.9 × 10 ¹¹	Q	1.048
1313–333	J1316–3338	[HB89] 1313–333	1FGL J1316.1–3341	321	8.1 ± 1.1	2.34 ± 0.08	(4.6 ± 0.7) × 10 ⁴⁷	1.33 ± 0.21	3.7 × 10 ¹²	Q	1.210
1322–428	J1325–4301	Cen A	1FGL J1325.6–4300	628	20.4 ± 1.5	2.71 ± 0.06	(1.3 ± 0.1) × 10 ⁴¹	0.72 ± 0.11	8.6 × 10 ¹⁰	G	0.002
1323–526	J1326–5256	PMN J1326–5256	1FGL J1327.0–5257	387	11.4 ± 1.3	2.33 ± 0.07		1.19 ± 0.18	2.7 × 10 ¹¹ *	B	>0.24
1325–558	J1329–5608	PMN J1329–5608	1FGL J1329.2–5605	244	13.9 ± 1.6	2.56 ± 0.08		0.47 ± 0.08	5.6 × 10 ¹⁰ *	B	>0.13
1344–376	J1347–3750	PMN J1347–3750	1FGL J1347.8–3751	73	5.0 ± 1.1	2.70 ± 0.15	(3.7 ± 0.8) × 10 ⁴⁷	0.20 ± 0.04	1.2 × 10 ¹¹	Q	1.300
1424–418	J1427–4206	[HB89] 1424–418	1FGL J1428.2–4204	241	6.7 ± 0.9	2.31 ± 0.08	(5.7 ± 0.9) × 10 ⁴⁷	1.47 ± 0.23	1.1 × 10 ¹²	Q	1.522
1440–389	J1443–3908	PKS 1440–389	1FGL J1444.0–3906	289	2.3 ± 0.5	1.83 ± 0.08	(3.4 ± 0.7) × 10 ⁴⁴	0.09 ± 0.02	8.8 × 10 ⁹	B	0.065
1454–354	J1457–3539	PKS 1454–354	1FGL J1457.5–3540	3707	33.3 ± 1.2	2.27 ± 0.03	(3.0 ± 0.2) × 10 ⁴⁸	0.70 ± 0.11	1.7 × 10 ¹¹	Q	1.424
1501–343	J1505–3432	PMN J1505–3432	1FGL J1505.1–3435	49	1.1 ± 0.6	2.02 ± 0.19	(2.3 ± 1.4) × 10 ⁴⁷	0.14 ± 0.03	1.3 × 10 ¹¹	B	1.554
1600–489	J1603–4904	PMN J1603–4904	1FGL J1603.8–4903	737	19.0 ± 1.8	2.12 ± 0.05	(1.1 ± 0.1) × 10 ⁴⁶	0.22 ± 0.04	1.1 × 10 ¹⁰		0.180
1600–445	J1604–4441	PMN J1604–4441	1FGL J1604.7–4443	509	21.9 ± 1.8	2.46 ± 0.05		0.95 ± 0.15	7.2 × 10 ¹⁰ *		>0.01
1606–667	J1610–6649	PMN J1610–6649	1FGL J1610.6–6649	266	1.2 ± 0.3	1.60 ± 0.10		0.04 ± 0.01	>2.1 × 10 ¹¹ *		
1613–586	J1617–5848	PMN J1617–5848	1FGL J1617.7–5843	118	11.2 ± 1.7	2.72 ± 0.11		1.59 ± 0.24	7.5 × 10 ¹⁰ *	Q	
1610–771	J1617–7717	[HB89] 1610–771	1FGL J1617.9–7716	339	10.7 ± 1.2	2.53 ± 0.08	(1.5 ± 0.2) × 10 ⁴⁸	1.38 ± 0.21	9.6 × 10 ¹⁰	Q	1.710
1646–506	J1650–5044	PMN J1650–5044	1FGL J1650.4–5042	252	18.4 ± 2.2	2.55 ± 0.07					>0.09
1714–336	J1717–3342	PMN J1717–3342	1FGL J1717.9–3343	209	22.2 ± 2.8	2.42 ± 0.06		0.54 ± 0.09	1.2 × 10 ⁹ *	B	
1759–396	J1802–3940	PMN J1802–3940	1FGL J1802.5–3939	1207	19.6 ± 1.2	2.25 ± 0.04	(1.3 ± 0.1) × 10 ⁴⁸	0.89 ± 0.14	3.7 × 10 ¹¹	Q	1.320
1933–400	J1937–3958	PKS 1933–400	1FGL J1938.2–3957	145	6.0 ± 0.9	2.45 ± 0.09	(1.6 ± 0.3) × 10 ⁴⁷	1.40 ± 0.21	2.5 × 10 ¹³	Q	0.965
1954–388	J1957–3845	[HB89] 1954–388	1FGL J1958.4–3847	273	7.6 ± 0.9	2.47 ± 0.08	(6.8 ± 0.8) × 10 ⁴⁶	1.57 ± 0.24	2.2 × 10 ¹¹	Q	0.630
2005–489	J2009–4849	[HB89] 2005–489	1FGL J2009.5–4849	533	3.9 ± 0.5	1.90 ± 0.07	(6.4 ± 0.8) × 10 ⁴⁴	0.76 ± 0.12	2.5 × 10 ¹⁰	B	0.071
2052–474	J2056–4714	[HB89] 2052–474	1FGL J2056.3–4714	1052	17.0 ± 1.1	2.54 ± 0.05	(1.6 ± 0.2) × 10 ⁴⁸	1.74 ± 0.27	1.5 × 10 ¹²	Q	1.489
2136–428	J2139–4235	PMN J2139–4235	1FGL J2139.3–4235	1263	9.0 ± 0.7	2.08 ± 0.05				B	
2155–304	J2158–3013	[HB89] 2155–304	1FGL J2158.8–3013	7310	21.4 ± 0.8	1.91 ± 0.02	(9.7 ± 0.5) × 10 ⁴⁵	0.38 ± 0.06	7.7 × 10 ¹⁰	B	0.116
2204–540	J2207–5346	[HB89] 2204–540	1FGL J2207.8–5344	186	5.2 ± 0.8	2.50 ± 0.10	(2.2 ± 0.4) × 10 ⁴⁷	0.71 ± 0.11	2.2 × 10 ¹¹	Q	1.206

^a name in B1950.0 IAU format; ^b brightness temperatures available in the IFGL period; for sources without measured redshift a lower limit based on $z = 0$, or a lower redshift limit where available, is indicated with a star (*); ^c classifications as described in Sect. 2.1: Q: Quasar, B: BL Lac, G: galaxy

Table 2. Properties of additional γ -ray detections of TANAMI sources in the 1FGL period.

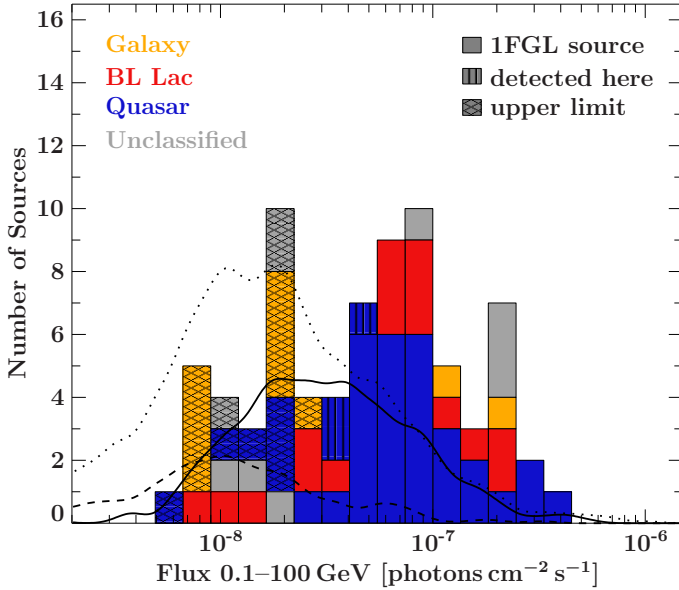
Name ^a	Source	Common name	TS	F (0.1–100 GeV) [10^{-8} ph cm $^{-2}$ s $^{-1}$]	LAT Γ	L (0.1–100 GeV) [erg s $^{-1}$]	S_{core} (8.4 GHz) [Jy]	T_B^b [K]	ID ^c	z	Sep. ^d [degree]	Conf95 $_{\gamma}^e$ [degree]
1505–496	J1508–4953	PMN J1508–4953	68	4.0 ± 0.9	2.18 ± 0.09	$(9.8 \pm 1.8) \times 10^{46}$	0.50 ± 0.08	2.2×10^{12}	Q	0.776	0.89	0.05
2149–306	J2151–3027	PKS 2149–306	61	5.1 ± 0.9	2.99 ± 0.16	$(2.2 \pm 0.5) \times 10^{48}$	1.27 ± 0.20	1.4×10^{12}	Q	2.345	1.51	0.02
2326–477	J2329–4730	[HB89] 2326–477	27	3.0 ± 1.0	3.0 ± 0.4	$(2.2 \pm 0.9) \times 10^{47}$	0.81 ± 0.13	9.7×10^{10}	Q	1.299	0.32	0.16

^a name in B1950.0 IAU format; ^b brightness temperatures available in the 1FGL period; for sources without measured redshift a lower limit based on $z = 0$ is indicated with a star (*); ^c classifications as described in Sect. 2.1: Q: Quasar, B: BL Lac, G: galaxy; ^d angular separation of the radio position of the TANAMI source and the closest 1FGL source; ^e semimajor axis of the 95% confidence region of the position of the closest 1FGL source

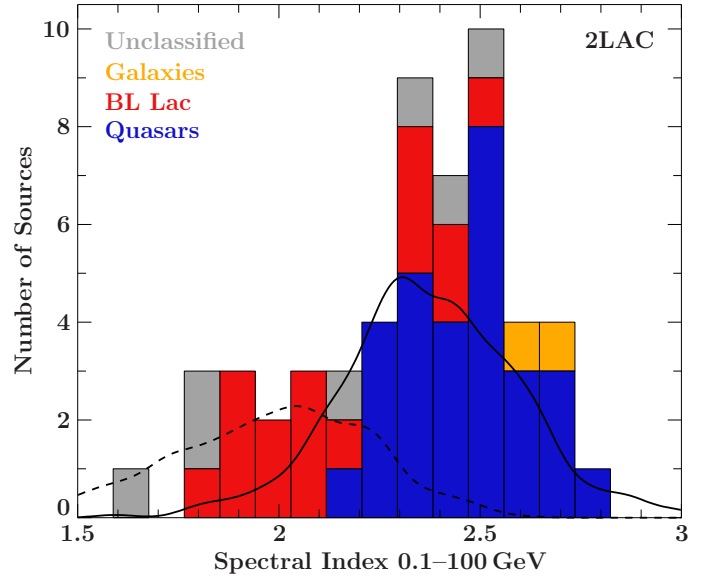
Table 3. Upper limits on the γ -ray emission of TANAMI sources not associated with 1FGL sources.

Name ^a	Source	Common name	TS	F (0.1–100 GeV) [10^{-8} ph cm $^{-2}$ s $^{-1}$]	L (0.1–100 GeV) [erg s $^{-1}$]	S_{core} (8.4 GHz) [Jy]	T_B^b [K]	ID ^c	z	Sep. ^d [degree]	Conf95 $_{\gamma}^e$ [degree]
0438–436	J0440–4333	[HB89] 0438–436	2	≤ 1.1	$\leq 6.8 \times 10^{47}$	0.59 ± 0.09	1.4×10^{11}	Q	2.863	1.69	0.02
0518–458	J0519–4546	PICTOR A	11	≤ 1.8	$\leq 2.7 \times 10^{43}$	0.55 ± 0.09	2.7×10^{10}	G	0.035	1.86	0.25
0527–359	J0529–3555	PMN J0529–3555	0	≤ 1.1	$\leq 2.1 \times 10^{45}$				0.323	1.51	0.10
1104–445	J1107–4449	[HB89] 1104–445	2	≤ 1.4	$\leq 1.7 \times 10^{47}$	1.41 ± 0.22	2.7×10^{11}	Q	1.598	0.77	0.21
1257–326	J1300–3253	PKS 1257–326	2	≤ 2.1	$\leq 1.4 \times 10^{47}$	0.13 ± 0.02	3.9×10^{10}	Q	1.256	0.35	0.13
1258–321	J1301–3226	ESO 443- G 024	18	≤ 2.4	$\leq 8.3 \times 10^{42}$			G	0.017	0.56	0.13
1333–337	J1336–3357	IC 4296	0	≤ 0.9	$\leq 1.5 \times 10^{42}$	0.22 ± 0.04	6.2×10^9	G	0.012	4.27	0.07
1549–790	J1556–7914	PKS 1549–79	5	≤ 2.2	$\leq 7.3 \times 10^{44}$	0.36 ± 0.06	1.6×10^{10}	G	0.150	2.23	0.09
1718–649	J1723–6500	NGC 6328	4	≤ 0.8	$\leq 9.3 \times 10^{41}$			G	0.010	3.57	0.17
1716–771	J1723–7713	PKS 1716–771	9	≤ 2.0		0.43 ± 0.07	9.0×10^{10} *			1.45	0.08
1733–565	J1737–5634	PKS 1733–56	0	≤ 0.8	$\leq 1.0 \times 10^{44}$	0.18 ± 0.03	4.9×10^{10}	G	0.098	4.83	0.10
1804–502	J1808–5011	PMN J1808–5011	23	≤ 2.2	$\leq 2.9 \times 10^{47}$	0.45 ± 0.07	5.9×10^{12}	Q	1.606	2.28	0.15
1814–637	J1819–6345	PKS 1814–63	5	≤ 2.1	$\leq 1.1 \times 10^{44}$	0.32 ± 0.05	6.5×10^{10}	G	0.063	1.95	0.18
1934–638	J1939–6342	PKS 1934–63	0	≤ 0.9	$\leq 4.5 \times 10^{44}$			G	0.180	5.74	0.40
2027–308	J2030–3039	PKS 2027–308	5	≤ 1.8	$\leq 1.3 \times 10^{46}$	0.08 ± 0.02	4.7×10^9		0.539	2.09	0.14
2106–413	J2109–4110	[HB89] 2106–413	0	≤ 0.6	$\leq 2.4 \times 10^{46}$	1.04 ± 0.16	1.6×10^{11}	Q	1.058	2.94	0.12
2152–699	J2157–6941	ESO 075- G 041	5	≤ 1.7	$\leq 1.6 \times 10^{43}$	0.43 ± 0.07	4.1×10^{10}	G	0.028	2.45	0.21
2355–534	J2357–5311	[HB89] 2355–534	2	≤ 1.7	$\leq 6.2 \times 10^{46}$	1.53 ± 0.23	5.0×10^{11}	Q	1.006	0.40	0.28

^a name in B1950.0 IAU format; ^b brightness temperatures available in the 1FGL period; for sources without measured redshift a lower limit based on $z = 0$ is indicated with a star (*); ^c classifications as described in Sect. 2.1: Q: Quasar, B: BL Lac, G: galaxy; ^d angular separation of the radio position of the TANAMI source and the closest 1FGL source; ^e semimajor axis of the 95% confidence region of the position of the closest 1FGL source

**Fig. 2.** Flux distribution of the *Fermi*/LAT-detected sources. For comparison the flux distributions of all 2LAC AGN (dotted line), the 2LAC quasars (solid line) and the 2LAC BL Lacs (dashed line) are shown as in Fig. 1.

of BL Lac objects have the same distribution as those of quasars. This result is consistent with previous studies, in which a relation between the gamma-ray spectral index and the peak frequency of the synchrotron component in the spectral energy distribution

**Fig. 3.** Distribution of the γ -ray spectral indices Γ of the *Fermi*/LAT-detected sources. The lines indicate the distributions of spectral indices of quasars (solid line) and BL Lac objects (dashed line) in the 2LAC (using a KDE as in Fig. 1).

has been found. The BL Lac objects are generally categorized as low-, intermediate- and high-synchrotron peaked sources (LSP, ISP, and HSP), which exhibit softer gamma-ray spectra with decreasing peak frequency (see, e.g., Fig. 17 of Ackermann et al.,

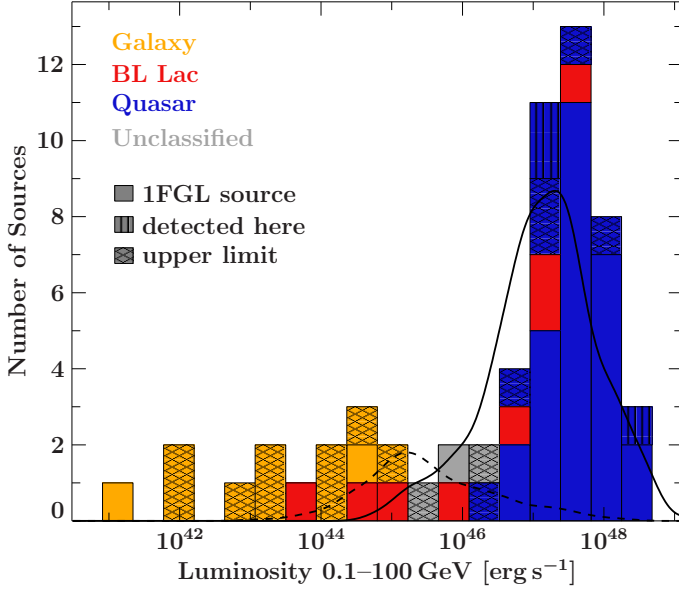


Fig. 4. Distribution of γ -ray luminosities. Radio galaxies have low luminosities, BL Lac objects are brighter and quasars the brightest objects in the sample. For comparison the luminosity distributions of quasars and BL Lac objects in the 2LAC are shown with the solid and dashed line, respectively (using a KDE as in Fig. 1).

2011b). This sequence extends to the quasars with the softest observed gamma-ray spectra.

The observed γ -ray luminosities (calculated as described in Sect. 2.4 and using a spectral index of $\Gamma = 2.4$ for the upper limits) show a clear dependence on the source classification (Fig. 4). Luminosities are low for radio galaxies, mostly intermediate for BL Lac objects, and high for quasars. The distribution of upper limits for the luminosity is consistent with that of the detected sources. The limits for the radio galaxies are clearly above the luminosity of Cen A ($(1.3 \pm 0.1) \times 10^{41} \text{ erg s}^{-1}$), which is the source with the lowest measured γ -ray luminosity in the sample. Upper limits for the quasars in the sample are not larger than the measured values but seem to have a similar distribution. A two-sample KS test does not indicate a significant statistical difference between both distributions (the probability that upper limits and luminosities of detected sources have the same underlying distribution is 19%).

The relation between the γ -ray luminosity and the spectral index is shown in Fig. 5 (this relation for all *Fermi*/LAT-detected AGN is shown in Fig. 24 of Abdo et al., 2010d). The Pearson correlation coefficient between both quantities (including the sources detected in the upper limit analysis) is 0.19. This value is comparable to the value of 0.17 found by Abdo et al. (2010d), who point out that the correlation might be influenced by instrumental detection limits and the Malmquist bias.

3.4. Radio-Gamma-ray-Relations

Using the results of the radio VLBI imaging, the core flux and the brightness temperature of the core for each TANAMI source were calculated (see Sect. 2.3). These radio properties were compared with the γ -ray properties. Figure 6 shows the relation of γ -ray flux and the 8.4 GHz radio core flux density. To quantify correlations we calculated Kendall's τ rank correlation coefficient for censored data (considering determined pairs only;

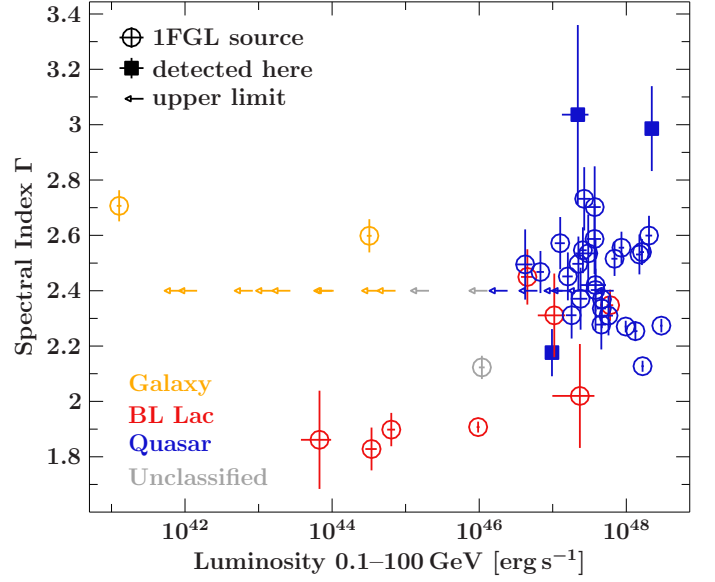


Fig. 5. Gamma-ray luminosities and spectral indices for different source classes. The upper limits are shown with the average index of $\Gamma = 2.4$.

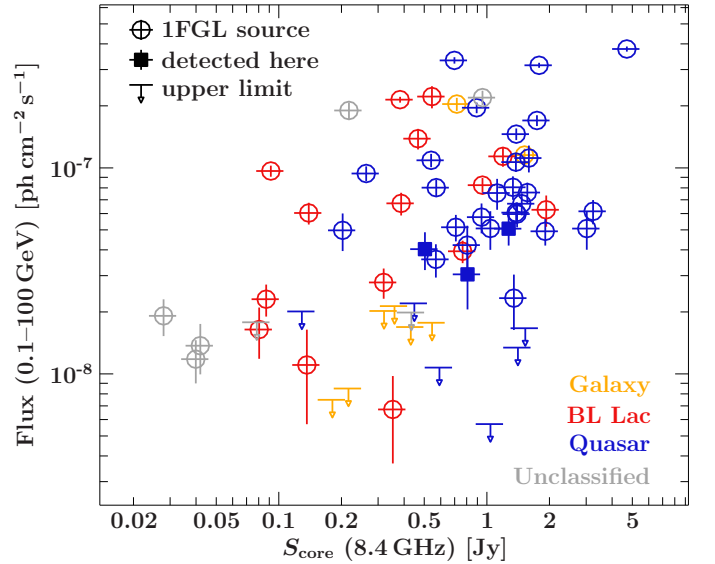


Fig. 6. Relation between γ -ray and radio core flux.

Helsel, 2005). Including the upper limits in this way the correlation coefficient is 0.29 with a p -value of 3.2×10^{-5} , confirming the correlation between these quantities. Using the distances and spectral indices of the sources the luminosities were calculated as described in Sect. 2.4. The relation between the radio and the γ -ray luminosity is shown in Fig. 7. Fitting a linear function to the logarithm of the luminosities yields $L_\gamma \propto L_r^{0.89 \pm 0.04}$. The clear correlation between the luminosities is induced by the source distances; thus the method of Akritas & Siebert (1996) is used, which yields a partial correlation coefficient of 0.30 with a p -value of 4.3×10^{-3} between radio and γ -ray luminosity given the redshift. These values mainly reflect the correlation between the fluxes. It is, however, necessary to consider that most of the sources show significant variability in both energy bands and that there might be source-dependent time delays between both bands. For that reason the obtained degree of correlation might

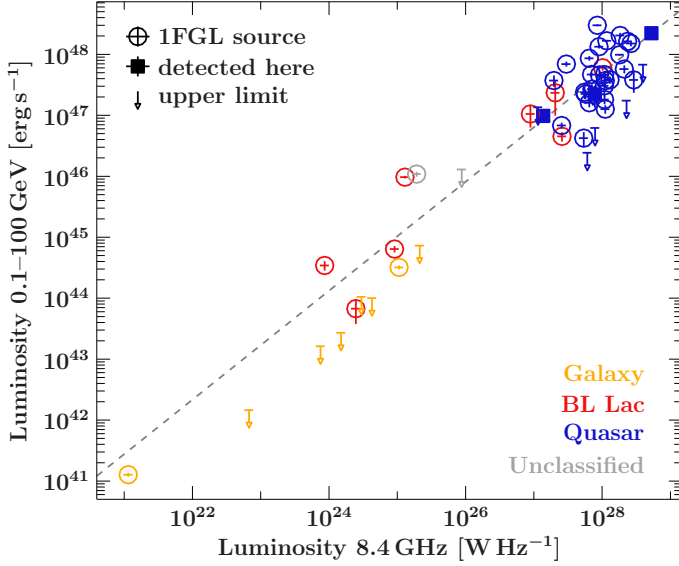


Fig. 7. Relation between γ -ray and radio core luminosity. The dashed line shows the fitted power law relation between both luminosities. Upper limits tend to be below this relation.

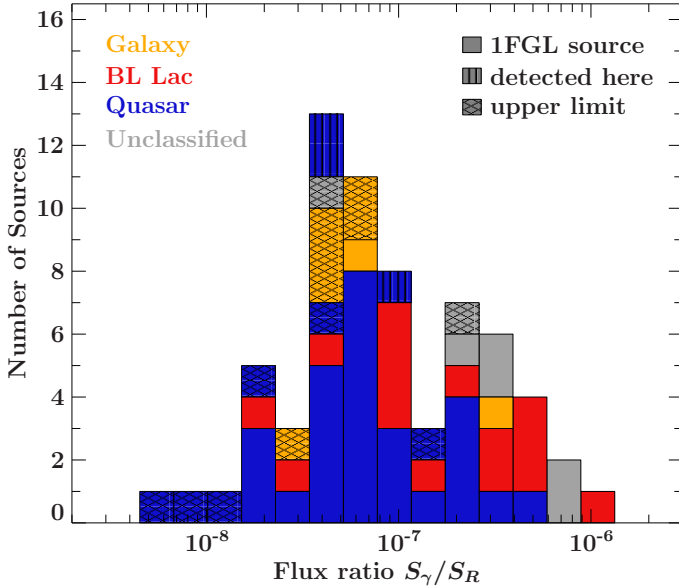


Fig. 8. Distribution of the ratio between γ -ray and radio core flux. The stacked histogram is color-coded according to the source class.

be decreased unless corresponding time periods are used for the observations in both energy bands.

Figures 6 and 7 show that the γ -ray/radio brightness ratio is larger for TANAMI sources that are detected with *Fermi* than for the *Fermi*/LAT-undetected sources in the sample, where the upper limit is used as the flux. A significant statistical difference between the distributions of the luminosity ratios of *Fermi*/LAT detected and undetected sources is found. A two-sample KS test yields a probability of only 1.2% that both ratios were drawn from the same underlying distribution. The γ -ray loudness, which is defined here simply as the ratio of the integral γ -ray flux to the radio flux density of the VLBI core, is presented in Fig. 8 for the AGN in the sample. The distribution of the γ -ray loudness indicates a similar dependence on the source class as the γ -ray spectral index distribution: while the quasars

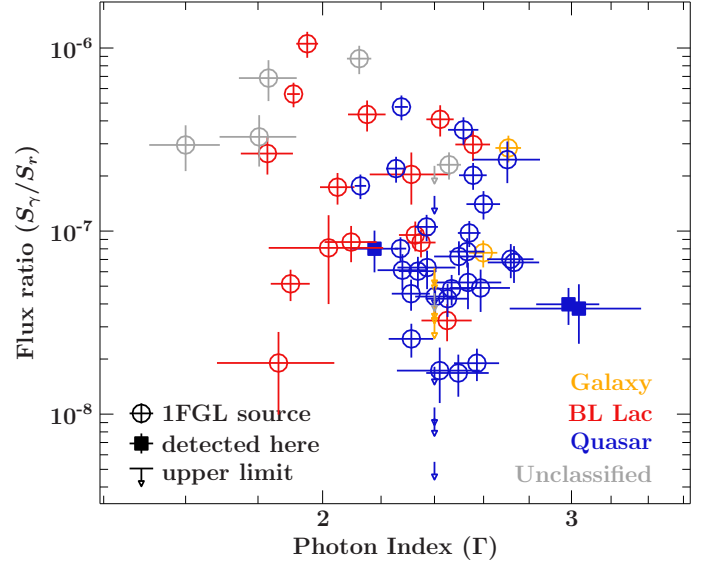


Fig. 9. Scatter plot of γ -ray loudness (defined here as a flux ratio) and γ -ray spectral index.

are less γ -ray loud, the BL Lac objects cover a broader range with a slightly higher averaged γ -ray loudness. This dependence is consistent with a shift of the peak frequencies in the SEDs. If the synchrotron peak is shifted to higher frequencies the flux density in the radio band decreases, whereas a shift of the high-energy peak in the SED towards higher frequencies increases the flux in the observed γ -ray band. At the same time the spectrum in this band hardens. This interpretation is strengthened by an observed anti-correlation between the γ -ray spectral index, Γ , and the γ -ray loudness as it is defined here (Fig. 9). We obtain a p -value of 1.0% for a Kendall τ rank correlation.

Figure 10 shows the distribution of the brightness temperatures. The distribution is consistent with that obtained from the first TANAMI observations shown by Ojha et al. (2010). A broadly similar distribution of brightness temperatures is also seen in the MOJAVE survey (Kovalev et al., 2009). The highest observed value is above 10^{13} K. Due to the theoretical limits on the brightness temperature in the source frame, such as the inverse Compton limit of $\sim 10^{12}$ K (Kellermann & Pauliny-Toth, 1969), the larger brightness temperatures observed here are a clear indication of strong Doppler boosting (as discussed, e.g., by Tingay et al., 2001b). The relation between brightness temperatures and the γ -ray luminosity is shown in Fig. 11. There is an indication that the brightness temperature of the radio core increases with increasing γ -ray luminosity. The partial correlation coefficient between γ -ray luminosity and brightness temperature, given the redshift, is 0.25 with a p -value of 1.8×10^{-3} . *Fermi*/LAT undetected sources, of which a large fraction have a low ratio of γ -ray to radio luminosity (Fig. 7), also tend to have lower brightness temperatures.

4. Individual Sources

In this section we comment on sources for which a classification has been added and/or the Véron-Cetty & Véron (2006) classification has been changed. Additionally, sources where our γ -ray analysis revealed new results are discussed.

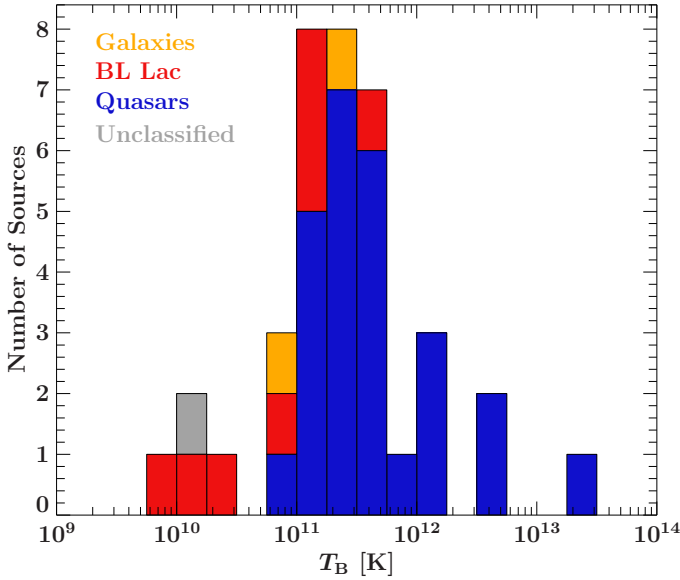


Fig. 10. Distribution of brightness temperatures in the TANAMI sample.

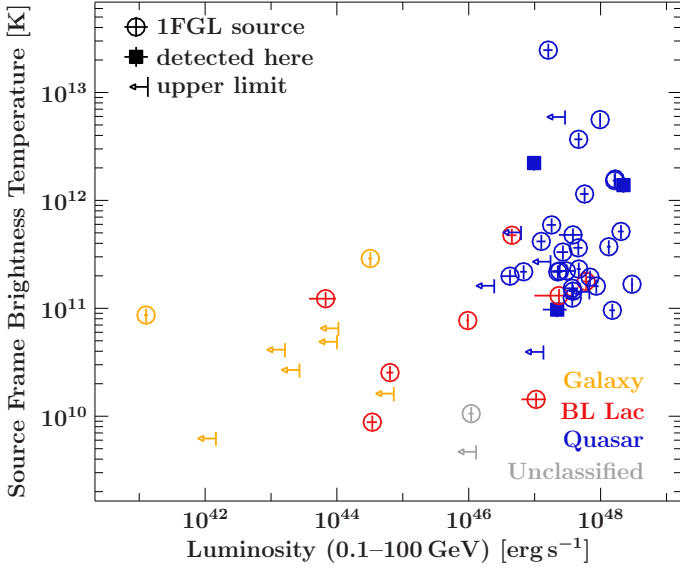


Fig. 11. Relation between γ -ray luminosity and brightness temperature of the radio core.

[HB89] 0208–512: We use a classification as quasar instead of the BL Lac classification because Wilkes (1986) find a MgII line with an equivalent width of 18 \AA .

PKS 0302–623: We use a quasar classification for this object, as it has properties of a flat spectrum radio quasar (see, e.g., Healey et al., 2008).

[HB89] 0332–403: The redshift of this source, which is by far the most luminous BL Lac in our sample, is difficult to determine. We use $z = 1.351$, which is based on a single weak Mg II emission line (Bergeron et al., 2011) and is consistent with the photometric redshift of this source (Rau et al., 2012, $z = 1.47^{+0.11}_{-0.12}$). If the source classification as a BL Lac is correct,

PKS 0332–403 would be one of only two BL Lac objects with $z > 1.3$ in the *Fermi* 2FGL sample (Rau et al., 2012).

We note, however, that the classification of this source as a BL Lac by Véron-Cetty & Véron (2006) seems not to be very secure. It seems to go back to Impey & Tapia (1990), who found that PKS 0332–403 has a high degree of polarization (14.7%), above their 2.5% for classifying sources as blazars, but little further work appears to have been done on classifying PKS 0332–403. The more specialized catalog of BL Lacs by Padovani & Giommi (1995) did not contain PKS 0332–403. The updated version of that catalog, Version 4.2 of the Roma-BZCAT catalog (Massaro et al., 2009), lists PKS 0332–403 as only a “BL Lac candidate”. Tornaiainen et al. (2008) and Tornikoski et al. (2001) are more conservative and list PKS 0332–403 as a highly polarized quasar and possibly a GPS source.

[HB89] 0438–436: This is a very bright and luminous radio quasar at high redshift ($z = 2.863$) that was not detected in the first 11 months of *Fermi*/LAT data. The upper limit on the luminosity, given by the flux upper limit and the distance, is, however, in the range of a relatively luminous quasar.

PKS 0447–439: We do not use a redshift for this source, because it is uncertain and some of its estimates are even contradictory, as discussed by Pita et al. (2014).

ESO 362-G021 (0521–365): Instead of a BL Lac classification we use a galaxy classification based on the presence of strong emission lines (see, e.g., Falomo et al., 1994).

PKS 0625–354 This source has a large-scale FR-I radio morphology, but its optical spectrum indicates a BL Lac classification (Wills et al., 2004), which we use in this work. PKS 0625–354 is one of five misaligned radio galaxies discussed in the 1LAC (Abdo et al., 2010d). The TANAMI mas-scale image shows a single-sided jet with a strong core component and is thus consistent with the inner parsec-scale jet being oriented close to the line of sight.

PKS 0745–330 has a nearby 1FGL source, but the radio position is slightly outside the positional 95% confidence region of the 1FGL source indicating that it might not be the correct counterpart. In the 2FGL catalog (2FGL; Nolan et al., 2012) the separation is, however, smaller and the radio and the 2FGL sources are associated. As the 1FGL and 2FGL sources are clearly associated, we use the 1FGL as counterpart for 0745–330 here. Testing the other case as well, in which the γ -ray source is not the correct counterpart for 0745–330, the modeling of a γ -ray source at the radio position in addition to the nearby 1FGL source yields a TS of 12.7 and a flux of $\leq 0.6 \times 10^{-7} \text{ ph cm}^{-2} \text{ s}^{-1}$.

PKS 1057–79: Sbarufatti et al. (2009), who measured the redshift of $z = 0.581$ of this source, list PKS 1057–79 as a BL Lac object but suggest classifying it as a broad-line AGN based on the observed emission lines. In our work we use the BL Lac classification. The γ -ray luminosity and the spectral index of this object are in between the typical values for quasars and BL Lacs.

ESO 443-G024 (1258–321): This is the brightest galaxy in the cluster ACO 3537 (Hudson et al., 2001), but has blazar properties as well. With a TS below 25 the source is not detected by *Fermi*/LAT. Modeling the source yields a TS of 18.2, a flux of $(1.6 \pm 0.8) \times 10^{-8} \text{ ph cm}^{-2} \text{ s}^{-1}$, and a spectral index 2.33 ± 0.22 .

PKS 1440–389: PKS 1440–389 has a very hard γ -ray spectrum with spectral index of 1.83 ± 0.08 . A BL Lac classification is likely (see, e.g., Jackson et al., 2002; Mao, 2011). The TANAMI VLBI image reveals a weak radio core with $S_{\text{Core}} \sim 50 \text{ mJy}$ and an extension to the southwest.

PMN J1508–4953: This source, which was not in 1FGL, was detected in our analysis. Modeling a γ -ray source at the radio position of that source yields a TS of 68, a flux of $(4.0 \pm 0.9) \times 10^{-8} \text{ ph cm}^{-2} \text{ s}^{-1}$, and a spectral index 2.18 ± 0.09 . These values are consistent with the source’s counterpart in the 2FGL catalog (2FGL J1508.5–4957). The properties obtained during the first two years of *Fermi* operation are TS = 70, a flux of $S_{\text{ph}} = (5.5 \pm 0.9) \times 10^{-8} \text{ ph cm}^{-2} \text{ s}^{-1}$ and a spectral index of 2.61 ± 0.09 .

PKS 2149–306: This source is clearly detected in the γ -ray data with a TS of 61, a flux of $(5.1 \pm 0.9) \times 10^{-8} \text{ ph cm}^{-2} \text{ s}^{-1}$, and a very soft spectral index of 2.99 ± 0.16 . With this value, the source has one of the steepest gamma-ray spectra of all AGN detected with *Fermi*/LAT. The source is included in the 2FGL catalog as 2FGL J2151.5–3021 and has the following properties averaged over the first two years of *Fermi*: TS = 168, $S_{\text{ph}} = (6.4 \pm 0.7) \times 10^{-8} \text{ ph cm}^{-2} \text{ s}^{-1}$, and a spectral index of 3.00 ± 0.09 .

[HB89] 2326–477: is detected with a TS of 27, a flux of $(3.0 \pm 1.0) \times 10^{-8} \text{ ph cm}^{-2} \text{ s}^{-1}$, and a spectral index of 3.0 ± 0.4 . The source is included in 2FGL as 2FGL J2329.7–4744 with TS = 25 and a flux of $S_{\text{ph}} = (1.6 \pm 0.7) \times 10^{-8} \text{ ph cm}^{-2} \text{ s}^{-1}$. The spectral index is 2.58 ± 0.22 . With a flux of $(7.7 \pm 2.0) \times 10^{-8} \text{ ph cm}^{-2} \text{ s}^{-1}$ this source was brightest in the tenth of the 24 monthly bins in the 2FGL catalog. The variability explains the slightly higher TS and flux average in the 1FGL period compared to the full 2FGL period.

5. Discussion

We present γ -ray properties of the TANAMI sample based on the data obtained with *Fermi*/LAT during its first 11 months of operation. A total of 54 out of the 75 AGN from this sample can be associated with γ -ray sources from the 1FGL catalog. All BL Lac objects (16/16) and a large fraction of quasars (29/38) are detected in the γ -ray regime, whereas from the 11 radio galaxies only the closest one (Cen A) and one with indications of being a BL Lac object (PKS 0521–365) were detected with *Fermi*/LAT. The low number of γ -ray detected radio galaxies is consistent with 1LAC and 2LAC where the dominating fraction of *Fermi*/LAT-detected AGN are blazars. In 2LAC the “clean” sample of 886 AGN includes 395 BL Lac objects and 310 quasars but only 8 misaligned AGN (Ackermann et al., 2011b). There does not appear to be a significant difference in the distribution of redshifts for *Fermi*/LAT detected and non-detected sources as was also found for the MOJAVE sample (Lister et al., 2011).

For the 21 AGN without γ -ray counterparts we presented upper limits on the γ -ray flux. In three cases (PKS 1505–496, PKS 2149–306, PKS 2326–477) the TS was high enough to indicate a detection. All of these three sources are included as detections in the 2FGL catalog. We note that the upper limits on γ -ray luminosity for quasars in the sample are comparable to the measured values, and both the limits and the values have a similar distribution. This suggests that the luminosity of the undetected quasars will be comparable to their upper limits. Alternatively, there could be a class of fainter (at γ -ray energies) quasars in the sample.

A weak correlation between radio and γ -ray fluxes, as well as an analogous partial correlation between the luminosities, of the sources in the TANAMI sample has been found in the first 11 months of *Fermi* science operations (see Figs. 6 and 7). Using EGRET data, Bloom (2008) found a correlation between the γ -ray and the radio luminosity at 8.4 GHz of $L_{\gamma} \propto L_r^{0.77 \pm 0.03}$. They could reproduce this relation using a synchrotron self-Compton model, but not with an external Compton model. Studies with *Fermi*/LAT data have shown that the γ -ray flux correlates well with compact (parsec scale) radio flux (Kovalev et al., 2009). For the TANAMI sample we find a relation of $L_{\gamma} \propto L_r^{0.89 \pm 0.04}$. A detailed interpretation of such a relation is difficult, because the emission in both bands is variable on different time scales and thus the ratio between the radio and γ -ray luminosity is not even constant for individual sources. With a longer set of contemporaneous observations, this problem could be addressed by searching for delays, or specifying well-justified averaging times. A perfect correlation between radio and γ -ray luminosity is, however, not expected, as it can be easily weakened, e.g., by different Doppler boosting in the radio and γ -ray regime. Different boosting between both energy regimes can for example originate from different Lorentz factors in the case of separated emission regions, but also from different spectral slopes in the radio and γ -ray regimes. For that reason, different angles to the line of sight of different sources in a sample decrease the observed correlation between the fluxes (e.g., Lister, 2007). Varying contributions of external Comptonization will scatter the radio γ -ray relation further. The fact that only weak correlations are found between the average radio and gamma-ray flux agrees with correlated variability being observed only rarely in both bands (see, e.g., Max-Moerbeck et al., 2013).

We find that BL Lacs in our sample tend to have harder spectra than quasars. Early *Fermi* results as well as the 2LAC show γ -ray loud quasars have soft spectra while the BL Lacs have a diverse range of spectral indices, where the gamma-ray spectra soften from HSP over ISP to LSP BL Lac objects (Ackermann et al., 2011b; Lott et al., 2012). The ratio of γ -ray and radio flux, which is used here to characterize the γ -ray loudness, is anti-correlated with the spectral index in the γ -ray band, which is consistent with the above-mentioned dependence on the synchrotron peak frequency.

Most γ -ray upper limits for undetected sources tend to be smaller than required for fitting to the average radio- γ -ray relation in the sample, i.e., many of the *Fermi*/LAT-undetected AGN are less γ -ray-loud than the detected sources. However, it has to be noted that the derived γ -ray loudness can be influenced by variability. As shown by Abdo et al. (2010d, their Figs. 11 and 19), most AGN show strong variability in the γ -ray regime. Flares can strongly increase the measured average flux. Additionally the sources are variable in the radio regime. Considering possible emission delays between the bands, the selection of corresponding time intervals is necessary for a better correlation. (Pushkarev et al., 2010, find delays of about

1.2 months in the source frame). Further temporal studies, including a search for delays and a comparison of the jet speeds with γ -ray properties will be carried out with more TANAMI epochs over a longer time period.

We find indications that γ -ray luminous AGN in the TANAMI sample have larger radio core brightness temperatures than γ -ray fainter sources. A comparison of brightness temperatures of strong EGRET sources and EGRET-undetected sources did not show this relation (Tingay et al., 1998, who used mainly observations at lower frequencies). A relation between brightness temperature and γ -ray brightness has been found in the MOJAVE sample (Kovalev et al., 2009; Lister et al., 2011).

6. Conclusions

The radio and γ -ray properties of the TANAMI AGN sample were investigated using data obtained during the first 11 months *Fermi* operations. Over 70% of the sample had already been detected by *Fermi*/LAT and the rates of detection for quasars and radio galaxies are consistent with that found in other samples and with studies using more γ -ray data.

For those TANAMI sources not in the 1FGL list, an upper limit analysis was performed. Three new γ -ray sources were significantly detected by this analysis. The luminosities of the rest of the undetected quasars are likely to be close to the upper limits reported here. The undetected sources have lower γ -ray-to-radio luminosity ratio and lower brightness temperatures which fits the picture of Doppler boosting playing a dominant role in determining the γ -ray state of an AGN.

A relation between γ -ray and radio flux was apparent and the brightness temperatures of radio cores were found to scale with the γ -ray luminosity. Some sources have brightness temperatures well above the inverse Compton limit suggesting strong Doppler boosting.

Similar studies of the TANAMI sample will be made using LAT data for different time ranges to tease out variations on different timescales. As enough epochs of data are now becoming available for most of the TANAMI sample, future studies will include VLBI kinematics. The TANAMI team has upcoming observations with Gemini South, which should allow optical identifications and redshift measurements for the large fraction of TANAMI sources (mostly fainter and poorly studied new *Fermi*/LAT detections) that do not have them, significantly improving our statistics.

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References

- Abdo A.A., Ackermann M., Ajello M., et al., 2010a, *ApJ* 719, 1433
- Abdo A.A., Ackermann M., Ajello M., et al., 2010b, *Science* 328, 725
- Abdo A.A., Ackermann M., Ajello M., et al., 2010c, *ApJS* 188, 405 (1FGL)
- Abdo A.A., Ackermann M., Ajello M., et al., 2010d, *ApJ* 715, 429 (1LAC)
- Ackermann M., Ajello M., Albert A., et al., 2012, *ApJS* 203, 4
- Ackermann M., Ajello M., Allafort A., et al., 2011a, *ApJ* 741, 30
- Ackermann M., Ajello M., Allafort A., et al., 2011b, *ApJ* 743, 171 (2LAC)
- Ackermann M., Ajello M., Atwood W.B., et al., 2015, *ApJ* in press (3LAC; arXiv:1501.06954)
- Akritas M.G., Siebert J., 1996, *MNRAS* 278, 919
- Atwood W.B., Abdo A.A., Ackermann M., et al., 2009, *ApJ* 697, 1071
- Bergeron J., Boissé P., Ménard B., 2011, *A&A* 525, A51
- Blandford R.D., Rees M.J., 1978, In: A. M. Wolfe (ed.) *BL Lac Objects.*, p.328
- Bloom S.D., 2008, *AJ* 136, 1533
- Cash W., 1979, *ApJ* 228, 939
- Cohen M.H., Lister M.L., Homan D.C., et al., 2007, *ApJ* 658, 232
- Danziger I.J., Fosbury R.A.E., Goss W.M., Ekers R.D., 1979, *MNRAS* 188, 415
- Falomo R., Scarpa R., Bersanelli M., 1994, *ApJS* 93, 125
- Fossati G., Maraschi L., Celotti A., et al., 1998, *MNRAS* 299, 433
- Ghisellini G., Maraschi L., Tavecchio F., 2009, *MNRAS* 396, L105
- Ghisellini G., Tavecchio F., Foschini L., et al., 2010, *MNRAS* 402, 497
- Hartman R.C., Bertsch D.L., Fichtel C.E., et al., 1992, *ApJ* 385, L1
- Healey S.E., Romani R.W., Cotter G., et al., 2008, *ApJS* 175, 97
- Helene O., 1983, *NIMPR* 212, 319
- Helsel D.R., 2005, *Nondetects and Data Analysis: Statistics for Censored Environmental Data*, Wiley, New York
- Hudson M.J., Lucey J.R., Smith R.J., et al., 2001, *MNRAS* 327, 265
- Impey C., 1996, *AJ* 112, 2667
- Impey C.D., Tapia S., 1990, *ApJ* 354, 124
- Jackson C.A., Wall J.V., Shaver P.A., et al., 2002, *A&A* 386, 97
- Kellermann K.I., Pauliny-Toth I.I.K., 1969, *ApJ* 155, L71
- Kovalev Y.Y., Aller H.D., Aller M.F., et al., 2009, *ApJL* 696, L17
- Kovalev Y.Y., Kellermann K.I., Lister M.L., et al., 2005, *AJ* 130, 2473
- León-Tavares J., Valtaoja E., Tornikoski M., et al., 2011, *A&A* 532, A146
- Linford J.D., Taylor G.B., Schinzel F.K., 2012, *ApJ* 757, 25
- Lister M.L., 2007, In: Ritz S., Michelson P., Meegan C.A. (eds.) *The First GLAST Symposium*, Vol. 921. AIPC, p.345
- Lister M.L., Aller M., Aller H., et al., 2011, *ApJ* 742, 27
- Lister M.L., Aller M.F., Aller H.D., et al., 2013, *AJ* 146, 120
- Lister M.L., Homan D.C., 2005, *AJ* 130, 1389
- Lister M.L., Homan D.C., Kadler M., et al., 2009, *ApJL* 696, L22
- Lott B., Cavazzuti E., Cutini S., et al., 2012, In: *Proc. “Fermi & Jansky”*. eConf C1111101
- Lovell J.E.J., McCallum J.N., Reid P.B., et al., 2013, *Journal of Geodesy* 87, 527
- Mao L.S., 2011, *New Astronomy* 16, 503
- Maraschi L., Ghisellini G., Celotti A., 1992, *ApJL* 397, L5
- Massaro E., Giommi P., Leto C., et al., 2009, *A&A* 495, 691
- Mattox J.R., Bertsch D.L., Chiang J., et al., 1996, *ApJ* 461, 396
- Max-Moerbeck W., Richards J.L., Pavlidou V., et al., 2013, In: *2012 Fermi Symposium proceedings*. eConf C121028
- Moellenbrock G.A., Fujisawa K., Preston R.A., et al., 1996, *AJ* 111, 2174
- Müller C., Kadler M., Ojha R., et al., 2014, *A&A* 569, A115
- Müller C., Kadler M., Ojha R., et al., 2011, *A&A* 530, L11
- Nolan P.L., Abdo A.A., Ackermann M., et al., 2012, *ApJS* 199, 31 (2FGL)
- Ojha R., Fey A.L., Johnston K.J., et al., 2004, *AJ* 127, 1977
- Ojha R., Kadler M., Böck M., et al., 2010, *A&A* 519, A45
- Padovani P., Giommi P., 1995, *MNRAS* 277, 1477
- Pita S., Goldoni P., Boisson C., et al., 2014, *A&A* 565, A12
- Pushkarev A.B., Kovalev Y.Y., Lister M.L., 2010, *ApJ* 722, L7

- Rau A., Schady P., Greiner J., et al., 2012, *A&A* 538, A26
- Sbarufatti B., Ciprini S., Kotilainen J., et al., 2009, *AJ* 137, 337
- Shaw M.S., Filippenko A.V., Romani R.W., et al., 2013, *AJ* 146, 127
- Shaw M.S., Romani R.W., Cotter G., et al., 2012, *ApJ* 748, 49
- Tavecchio F., Ghisellini G., Bonnoli G., Ghirlanda G., 2010, *MNRAS* 405, L94
- Taylor G.B., Healey S.E., Helmboldt J.F., et al., 2007, *ApJ* 671, 1355
- Thompson D.J., Bertsch D.L., Fichtel C.E., et al., 1993, *ApJS* 86, 629
- Tingay S.J., Edwards P.G., 2002, *AJ* 124, 652
- Tingay S.J., Jauncey D.L., King E.A., et al., 2003, *PASJ* 55, 351
- Tingay S.J., Murphy D.W., Lovell J.E.J., et al., 1998, *ApJ* 497, 594
- Tingay S.J., Preston R.A., Jauncey D.L., 2001a, *AJ* 122, 1697
- Tingay S.J., Preston R.A., Lister M.L., et al., 2001b, *ApJ* 549, L55
- Tornainen I., Tornikoski M., Turunen M., et al., 2008, *A&A* 482, 483
- Tornikoski M., Jussila I., Johansson P., et al., 2001, *AJ* 121, 1306
- Tzioumis A.K., Tingay S.J., Stansby B., et al., 2010, *AJ* 140, 1506
- Urry C.M., Padovani P., 1995, *PASP* 107, 803
- Véron-Cetty M.P., Véron P., 2006, *A&A* 455, 773
- Wilkes B.J., 1986, *MNRAS* 218, 331
- Wills K.A., Morganti R., Tadhunter C.N., et al., 2004, *MNRAS* 347, 771
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